

IP Pegasi in outburst: Echelle spectroscopy & Modulation Doppler Tomography*

C. Papadaki¹, H.M.J. Boffin² and D. Steeghs³

(1) Vrije Universiteit Brussel, Pleinlaan 2, 1050 Brussels, Belgium

(2) ESO, Karl-Schwarzschild-Str. 2, 85748 Garching, Germany

(3) Department of Physics, University of Warwick, Coventry CV4 7AL, UK

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Abstract

We analyse a unique set of time-resolved echelle spectra of the dwarf nova IP Peg, obtained at the ESO NTT with EMMI on the 4th of August 1999, just one day after the peak of an outburst. The dataset covers the wavelength range of 4000–7500 Å and shows Balmer, HeI, HeII and heavier elements in emission. IP Peg was observed one day after the peak of an outburst.

The trailed spectra, spectrograms and Doppler maps show characteristics typical of IP Pegasi during the early stages of its outburst, such as prograde rotation from the accretion disc flow, chromospheric emission from the secondary and spiral arms. The high-ionisation line of HeII λ 4686 is the most centrally located line and its formation region has the greatest radial extension compared to the HeI lines. The Balmer lines extend from close to the white dwarf up to $\approx 0.45R_L$, with the outer radius gradually increasing when moving from H δ to H α . The application, for the first time, of the modulation Doppler tomography technique, maps any harmonically varying components present in the system configuration, and this variability information is not considered in standard Doppler tomography studies. We find, as expected, that part of the strong secondary star emission in Balmer and HeI lines is modulated predominantly with the cosine term, consistent with the emission originating from the irradiated front side of the mass-donor star, facing the accreting white dwarf. For the Balmer lines the level of the modulation, compared to the average emission, decreases when moving to higher series. Emission from the extended

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accretion disk appears to be only weakly modulated, with amplitudes of at most a few % of the non-varying disk emission. We find no evidence of modulated emission in the spiral arms, which if present, is relatively weak at that our signal-to-noise ratio was good enough to put an upper detection limit of any modulated emission at 5–6%. Only in one arm of the HeII $\lambda 4686$ line, is there a possibility of modulated emission, but again, we cannot be sure this is not caused by blending with the nearby Bowen complex of lines.

1 Introduction

Cataclysmic variable stars (CVs) consist of a Roche-lobe filling main sequence star, called the secondary star, orbiting a white dwarf - the primary star. Mass flows from the secondary star through the inner Lagrangian point towards the primary star, ultimately leading to the formation of an accretion disc (AD). After the AD's formation, a bright spot (BS) is formed at the contact point where the gas stream hits the AD. Dwarf novae (DNe) are a subclass of cataclysmic variables with two different brightness states, quiescence and outburst. In the latter, the system usually brightens by 2–5 mag with a time-scale that can vary from tens to hundreds of days.

IP Peg is the brightest eclipsing ($i \approx 81^\circ$) DN above the period gap, and was discovered by Lipovetsky & Stepanian (1981). It has a 3.8-h orbital period and a V magnitude in the range of 14–12 mag, the extreme values corresponding to quiescence and outburst, respectively.

From the first photometric light curves during quiescence and return to quiescence (Wood & Crawford, 1986; Wood et al., 1989), it was evident that the BS produces a prominent orbital hump, when it rotates into view from behind the disc with a single-stepped ingress and double-stepped egress superposed on a more gradual disc-eclipse. Infrared light curves showed 0.2 mag ellipsoidal variations from the secondary star superposed on a deep eclipse of the WD and the BS (Szkody & Mateo, 1986). Wood et al. (1989), based on the analysis of 39 eclipses, found out that IP Peg has a rapidly increasing orbital period. However, a new study performed by Wolf et al. (1993) and combining 49 eclipses along with previously reported ones, reached the conclusion that the orbital period varies sinusoidally with a period of 4.7 years. This variation they attributed to a third body, a late M dwarf star with a most probable mass of $0.10 M_\odot$. They also showed that the rate of shrinkage in disc radius during decline from outburst is constant during different outburst cycles, following an exponential decay. A study of the flickering in IP Peg (Bruch, 2000) showed that the BS is the dominant source of flickering, with the AD and WD contributing practically nothing.

Spectroscopy of the red star (Martin et al., 1989) showed TiO bands and the NaI doublet (red dwarf features), several broad double-peaked emission lines (disc features) and a narrow third peak in the emission lines moving in anti-phase with the double peaks (a chromospheric emission feature). Beekman et al. (2000) performed a TiO study of IP Peg and concluded that TiO instead of the NaI doublet should be

used for radial velocity studies of the secondary star. In this way they found that the contamination from disc emission, which in other studies returned an elliptical orbit, could be avoided and higher signal-to-noise ratio could be achieved. They gave a final projected orbital velocity of the companion star $K_2 = 331.3 \pm 5.8 \text{ km s}^{-1}$ and refined values of the primary and secondary star masses $M_1 = 1.05_{-0.07}^{+0.14} M_\odot$ and $M_2 = 0.33_{-0.05}^{+0.14} M_\odot$, respectively. Based on previous work and covering a complete cycle from quiescence to outburst and quiescence again, the spectral characteristics of IP Peg will now be described.

Quiescence spectra (Marsh, 1988) show broad double-peaked emission lines on a blue continuum, a strong orbital hump in the continuum (but much weaker in the lines) and a flat Balmer decrement indicative of optically thick emission. The trailed spectra show the signature of a prograde rotation of the AD, where the blue-shifted emission is eclipsed prior to the red-shifted one, a weak S-wave component in the red-shifted wing and a 10% blue/red asymmetry before eclipse, most probably arising from line emission from the BS.

Early rise-to-outburst spectra (Wolf et al., 1998) show Balmer emission, TiO absorption bands, KI and NaI absorption doublets, a weak S-wave and two bright spots in both Doppler and eclipse maps. The study of Wolf et al. (1998), as well as a re-analysis of their results using an extension of the classical eclipse mapping method (Bobinger et al., 1999), concluded that the first spot should be identified as the BS while two possibilities existed for the second spot. It either was the beginning of formation of a spiral arm or a second inner BS due to gas-stream overflow as predicted by, e.g., simulations performed by Armitage & Livio (1996).

Rise-to-outburst spectra (Harlaftis et al., 1994) show a decrease in the equivalent width of $H\alpha$ emission line, down to a level similar to that observed immediately after outburst and, simultaneously, weak $H\alpha$ emission from the secondary due to irradiation of the secondary star by the boundary layer.

Near-outburst spectra (Piche & Szkody, 1989) show unusually strong high-excitation lines, Balmer lines in emission, indications of a very disturbed AD with a non-uniform He emission distribution, an outflowing wind from the WD vicinity and chromospheric emission from the secondary.

Outburst spectra as presented by Marsh & Horne (1990) show: i) strong Balmer and HeII $\lambda 4686$ double-peaked emission lines rising from the AD, ii) non-Keplerian emission filling in the centre of HeII $\lambda 4686$, probably due to a compact outflowing wind or inflowing magnetic accretion columns from close to the WD, iii) Balmer but not HeII emission from the secondary, located in the polar region on the side facing the AD, iv) no evidence for an enhanced gas-stream compared to quiescence, v) large azimuthal asymmetry in the outer disc, with the blue-shifted side of the disc being brighter and possibly more extended at phase 0 and vi) a sharp transient component in the HeII line possibly generated from the interaction of the gas stream and disc. Marsh & Horne (1990) suggested that He emission can be explained as photoionisation of the disc by the boundary layer and found that Balmer emission in outburst can be attributed by 3/4 to the disc, 1/4 to irradiation of the red star and also to boundary layer irradiation. Another outburst study performed by

Steeghs et al. (1996) showed, from trailed spectra and Doppler maps, asymmetric disc emission, strong secondary star emission and a low velocity emission in the $H\alpha$ line only. This stationary emission, appearing as an emission blob in the centre of the Doppler maps, was interpreted as slingshot prominences, material trapped in magnetic loops and co-rotating with the secondary star. During another outburst maximum, Harlaftis et al. (1999) found metal line emission such as $MgII$. Doppler tomography of the He and metal lines locates those lines' emission on the inner Roche lobe of the secondary star.

Early decline-from-outburst spectra (Hessman, 1989) reveal for the first time in such strength a blend of high-excitation lines of $HeII$, $CIII$ and $NIII$ along with a very strong chromospheric emission-line component from the secondary.

Decline-from-outburst spectra (Morales-Rueda et al., 2000) show irradiation-induced emission from the companion star in Balmer, HeI , $HeII$, $MgII$ and CII lines. This emission, located only near the poles of the secondary, suggests shielding of the equatorial regions of the secondary star by the vertically extended AD.

Early-quiescence spectra in the infrared (Littlefair et al., 2001) reveal the existence of mirror eclipses in IP Peg. A mirror eclipse is an eclipse of the secondary star by an optically thin AD and appears in the trailed spectra as a reduction in the equivalent widths of the lines, at phase 0.5. A mirror eclipse also has a mirror-symmetry with the classical rotational disturbance, i.e. the red-shifted part of the line is eclipsed prior to the blue-shifted one.

Even though IP Peg had shown many peculiarities and novelties in the studies cited previously, it had to reveal another breakthrough. Steeghs et al. (1997) were the first to find convincing evidence for the existence of spiral structure in the AD of a close binary. Spiral waves had long been predicted in studies of tidal interactions and numerical simulations (see Boffin (2001) and references therein) and may play a crucial role in our understanding of the angular momentum transport. Steeghs et al. (1997) obtained spectrophotometric observations of IP Peg during the rise to an outburst, covering $H\alpha$ and HeI $\lambda 6678$. Disc emission was centred on the WD and had a strong azimuthal asymmetry in the forms of two spiral arms. The repeatability of the two-armed spiral structure was confirmed during another outburst maximum. Harlaftis et al. (1999) applied Doppler tomography of the $HeII$ $\lambda 4686$ line during outburst and clarified the shock nature of the spiral arms which, in the same map, showed a jump of a factor of more than 2 in intensity. Additionally they calculated an azimuthal extent of 90° , a 30° upper limit in opening angle and a 15% contribution of the shocks to the total disc emission. Morales-Rueda et al. (2000) showed that the spiral pattern persists even 5 and 6 days after the outburst maximum, without showing any diminution in strength. Time-resolved optical spectra obtained during late outburst by Baptista et al. (2000) reveal, through the eclipse mapping technique, that spirals are present even 8 days after the onset of an outburst. Their study showed a similar contribution to the continuum emission, as that measured during the peak of the outburst but also a retrograde azimuthal rotation of 58° compared to outburst maximum.

Baptista et al. (2005) performed a re-analysis of previous spectroscopic data of

IP Peg during outburst (Morales-Rueda et al., 2000), based on the eclipse mapping technique. The difference in their new study was based on the fact that they separated the inner and outer regions of the emission line light curves, obtaining in this way the “core” and “spiral arms” lights curves, respectively. In this way they improved the previously noisy view of the spiral arms and were able to trace them in detail. The two-armed spiral structures were clearly visible in all their eclipse maps, albeit the “blue” arm appearing farther out from the disc’s centre than the “red” one. However, their crucial findings were the sub-Keplerian velocities along the spiral structures, as well as the clear correlation between the opening angle of the spirals and the outburst stage (the later in the outburst, the smaller the opening angle of the spirals). Both these findings greatly favour the interpretation of the spiral asymmetric structures as tidally-induced spiral shocks instead of mere radiation patterns on an otherwise unperturbed AD.

A study performed by Steeghs (2001) presents a review of the spiral structure that far, and plots in detail the azimuthal angles and extent of the spirals for the HeII $\lambda 4686$ emission line during an outburst presented by Harlaftis (1999). In more detail, it is shown that the two spirals can be traced for almost 180° , while their velocity varies from 495 to 780 km s^{-1} . Under the assumption of Keplerian velocities, they show that the arms cover a substantial part of the disc, 0.3 – 0.9 times the distance to the inner Langragian point L_1 . It is pointed out that tracing the spiral arms and plotting their position and intensity as a function of azimuth throughout an outburst, or even better applying this procedure to a range of emission lines simultaneously, can yield important results on the evolution of the spirals. A future paper, as stated in the Discussion section, is intended to address this issue. A first approach on tracing the spiral arms of two different emission lines ($\text{H}\alpha$, HeI $\lambda 6678$), as well as plotting their emissivity as a function of azimuth, is presented in Steeghs (2001). It is shown that, even though there is a similar position for the spirals in terms of velocity as a function of azimuth, the intensity modulation is significantly different.

Our work is part of a project of studying CVs through high resolution time-resolved echelle spectra and applying the indirect imaging technique of Doppler tomography. The simultaneous study of several emission lines can yield important results on the AD structure, as presented for another DN, V2051 Oph, in Papadaki et al. (2008). Following the introduction on IP Peg one can conclude on the benefits of applying such an analysis to this object, especially since it was observed during outburst. In this way we can study the behaviour of the spiral arms and map the structure of the AD in detail. The application of Modulation Doppler tomography, an extension of the classic Doppler tomography method (Steeghs, 2003), gives us the opportunity to map emission sources which can vary harmonically with the orbital period. Such an analysis is a first for IP Peg.

This paper is organised as follows: in Sect. 2 we present our observations, in Sect. 3 the data reduction, in Sect. 4 we perform the spectral analysis and Sect. 5 deals with the application and results of Modulation Doppler tomography. Last, Sect. 6 is devoted to the discussion and summary of our results.

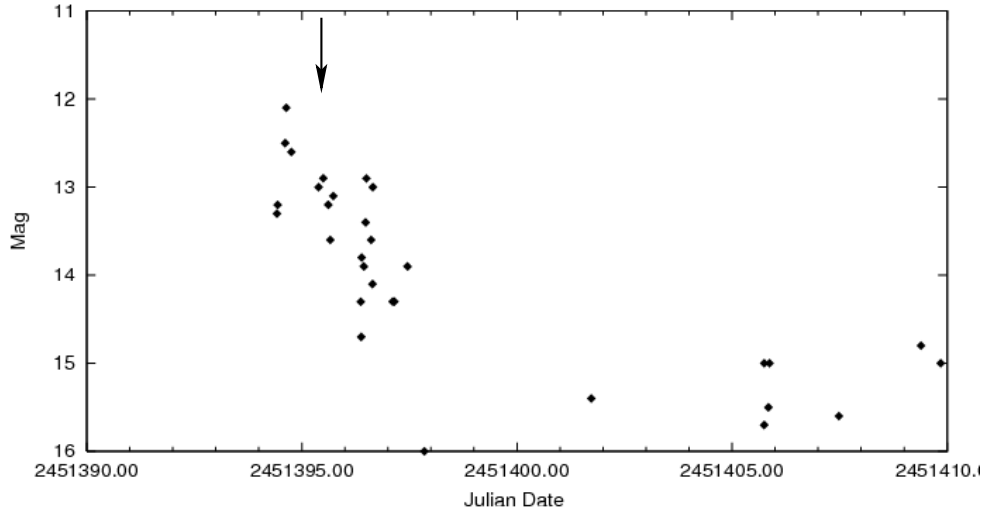


Figure 1: AAVSO long-term light curve of visual validated data. The arrow shows the timing of our observations, just 1 day after the peak of the outburst.

2 Observations

The observing run was conducted on the 4th of August 1999, at ESO’s NTT. Echelle spectroscopy was performed using the REMD mode of the EMMI spectrograph (Dekker et al., 1986). The echelle grating #9 and cross-disperser #3 resulted in 21 orders covering the wavelength range 4000–7500 Å. The dispersion increased between 0.09 and 0.17 Å pix⁻¹, from blue to red, respectively. IP Peg was observed for 5.7 h, covering 1.5 binary orbits. In total, 41 spectra with an exposure time of 300 sec each and a typical signal-to-noise ratio around 6, were obtained. The spectra are included in the data files available in electronic form as explained in Sect. 7.

Fig. 1 shows a long-term light curve of our system as generated from the AAVSO light curve generator tool¹, choosing only visual-validated data to appear, and a window of a few days around outburst. An inspection of the outburst maximum reveals that IP Peg was at the peak of its outburst just 1 day before our observing run, when it was also 1 magnitude brighter.

3 Data reduction

The resulting spectra were reduced using the standard echelle reduction procedures in IRAF¹ and the slit spectra reduction task `doecslit`. Whenever, both an inter-

¹<http://www.aavso.org/data/lcg/>

¹IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the

active and non-interactive mode existed for a task, the interactive one was always chosen.

Initially, the bias and lampflat frames were combined into one bias and flat frame with the `zerocombine` and `flatcombine` tasks, respectively. The spectra were then de-biased and a normalised flat was produced with the task `apflatten` in order to apply the flat-fielding process. `Apflatten`, after finding and tracing all apertures (defined as the sum of dispersion lines containing the signal of an order), also models both the profile and overall spectrum shape and removes it from the flat-field leaving only the sensitivity variations. The spectra were then divided by the normalised flat and thus flat-fielded. At this stage, the spectra were also corrected for cosmic rays with the use of the `cosmicrays` task, part of the `crutil` package.

The task `doecslit` connects and combines several individual tasks to provide a single, complete data reduction path. It provides a degree of guidance, automation and record keeping. Additionally, the user with the choice of the interactive mode (used in this set of spectra) and the use of a graphical interface, is able to be in control of the whole procedure, and if necessary, fine-tune the reduction process while it runs.

The actual extraction of the spectra was done by applying variance weighting, often called optimal extraction. For detailed information of this method the reader should refer to Horne (1986) and Marsh (1989), whose work was followed in IRAF. Dispersion correction was applied through the use of Thorium-Argon arc lamp spectra. The typical root mean square of the dispersion function that was fit to the total of identified lines in each order was $\approx 0.012 \text{ \AA}$. Last, the spectrophotometric standard LTT 7379 was used for the flux calibration of the spectra.

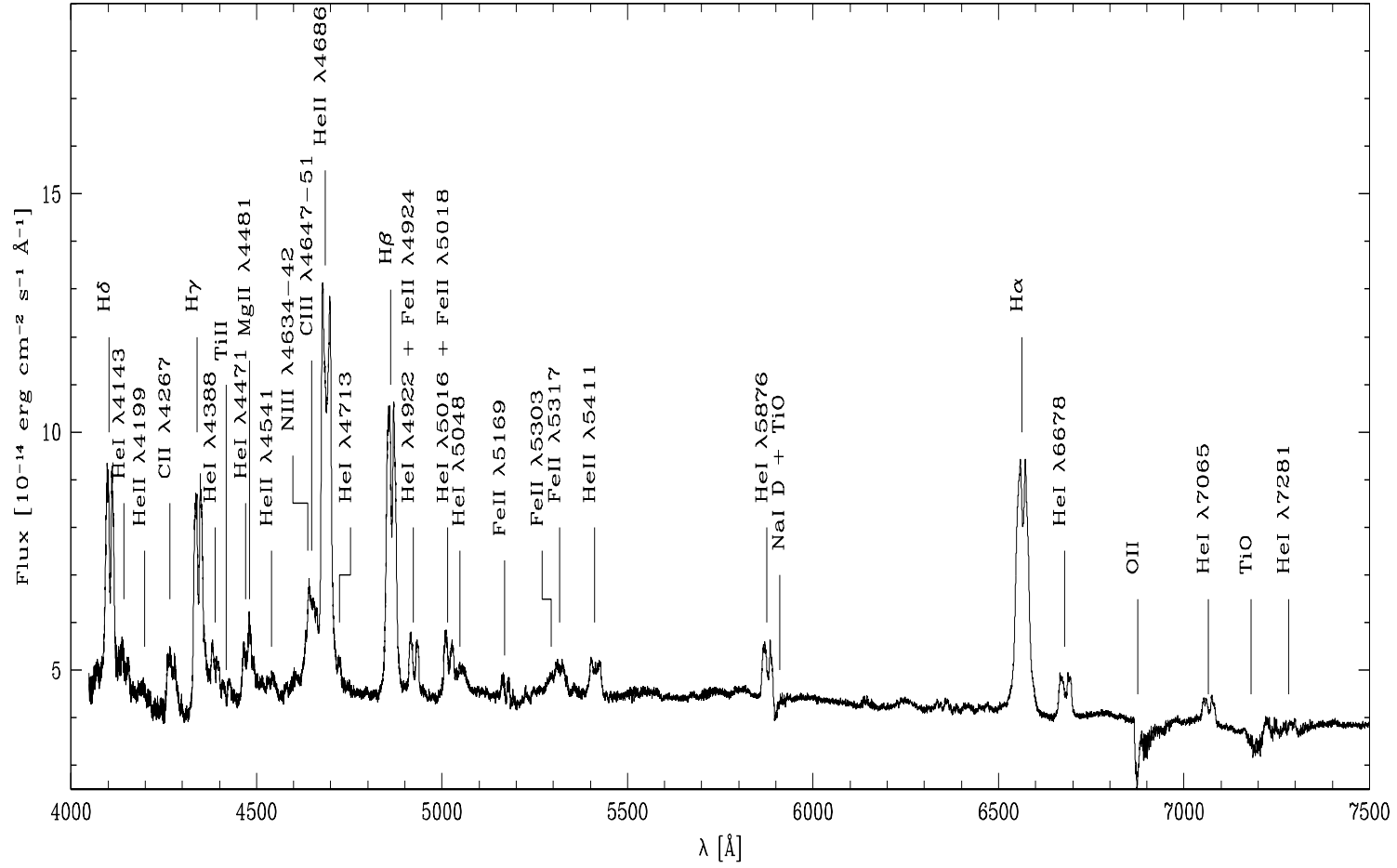


Figure 2: IP Peg average out-of-eclipse spectrum.

Table 1: Characteristics of emission lines.

	FWZI [km s ⁻¹]	EW [Å]	Integrated Flux [10 ⁻¹³ erg cm ⁻² s ⁻¹]
H δ	5695 $\pm$ 49	24.8 $\pm$ 0.9	11.1 $\pm$ 0.3
H γ	5945 $\pm$ 183	30.6 $\pm$ 1.3	12.5 $\pm$ 0.7
HeII λ 4686	5055 $\pm$ 68	51.3 $\pm$ 1.3	23.8 $\pm$ 0.7
H β	4237 $\pm$ 227	36.4 $\pm$ 0.9	16.3 $\pm$ 0.3
HeI λ 4922+FeII λ 4924	2101 $\pm$ 45	4.7 $\pm$ 0.1	2.2 $\pm$ 0.1
HeI λ 5876	1822 $\pm$ 34	5.5 $\pm$ 0.2	2.4 $\pm$ 0.1
H α	4023 $\pm$ 42	45.2 $\pm$ 0.7	18.8 $\pm$ 0.2
HeI λ 6678	2479 $\pm$ 74	6.7 $\pm$ 0.2	2.7 $\pm$ 0.1
HeI λ 7065	1818 $\pm$ 22	3.5 $\pm$ 0.1	1.4 $\pm$ 0.1

4 Spectral analysis

The average out-of-eclipse (phases 0.15–0.85, where phase 0 is that of mid-eclipse) spectrum is shown in Fig. 2. To check that spectra of incomplete orbits do not affect the line profiles of the resulting average out-of-eclipse spectrum we first computed and compared the average out-of-eclipse spectrum of the whole dataset (1.5 orbits) and that of one full orbit. The spectrum resulting from the whole dataset differed only in the Balmer lines which showed an enhancement of the blue-peak intensity. The rest of the lines showed minimal or no differences at all. Therefore, the out-of-eclipse spectra of one full orbit were used in Fig. 2. The orbital phases of the individual spectra were calculated according to the linear orbital ephemeris given by Wolf et al. (1993):

$$T_0[\text{HJD}] = 2445615.4224 + 0.^{\text{d}}15820616\text{E} \quad (1)$$

where T_0 is the time of mid-eclipse.

The outburst spectrum (Fig. 2) is characterised by several double-peaked emission lines, along with absorption features. There is strong Balmer and high-excitation emission from HeII λ 4686 and the CIII-NIII blend. Also present is weaker emission from HeI (λ 4143, 4388, 4471, 4713, 4922, 5016, 5048, 5876, 6678, 7065 and 7281), HeII (λ 4199, 4541 and 5411) and CII λ 4267 lines. IP Peg also shows emission from the heavier elements of FeII (FeII 42 triplet and FeII λ 5169, λ 5303), TiII λ 4418 and MgII λ 4481. Finally, there are the NaI doublet at 5890 and 5896 Å and TiO absorption bands at around 5890 and 7190 Å. Both features are characteristic of M-type stars, i.e. the same as the spectral type as the secondary star of IP Peg (Martin et al., 1987).

All emission lines appear with symmetric line wings, except the lines of the blue part of the spectrum where one cannot tell with certainty due to their severe blending with neighbouring lines. Concerning the blue-to-red peak strength, the Balmer lines seem to differentiate from the high-excitation lines. They have an

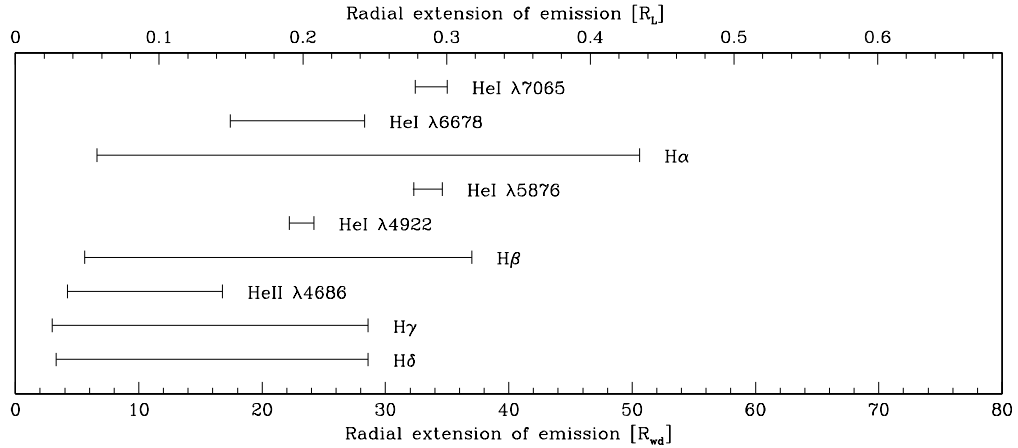


Figure 3: Radial extension of all emission lines. Bottom x-axis gives the radial extension as a factor of R_{wd} , while top x-axis gives it as a factor of the volume radius of the primary Roche lobe R_L

equal blue-to-red-peak strength (except H γ which has a stronger red peak) while the high excitation lines (blended or stand-alone) appear to have a stronger blue peak. For the HeI lines it is harder to distinguish between the strength of the peaks, as they are either located in the wings of stronger lines ($\lambda 4143$, $\lambda 4388$, $\lambda 4713$) or blended with equally strong lines ($\lambda 4471$, $\lambda 4922$, $\lambda 5016$, $\lambda 4143$). However, the stand-alone HeI lines ($\lambda 5876$, $\lambda 6678$, $\lambda 7065$) point to a similar behaviour with the Balmer lines.

In this study, we will focus on the 9 most prominent emission lines i.e. the Balmer lines, HeII $\lambda 4686$ and HeI ($\lambda 4922$, 5876 , 6678 , 7065). Table 1 shows the characteristics of the emission lines (full width at zero intensity -FWZI, equivalent width -EW, and integrated flux). However, caution should be drawn in the interpretation of these values for the severely blended lines, namely H δ , H γ , HeII $\lambda 4686$ and HeI $\lambda 5016$. The Balmer decrement from H α to H δ 1:0.77:0.68:0.55 shows a mix of optically thick and thin conditions.

Assuming a Keplerian velocity field and AD location of line emission, we can then estimate the radial extension of the emission lines given the fact that the velocities measured from the double-peak separation and FWZI are located in the outer and inner disc radius, respectively. Using equations (2) and (3) from Papadaki et al. (2008) and taking the tidal effect into account the resulting radial extension of the emission lines is shown in Fig. 3. The bottom x-axis is measured in terms of the WD radius R_{wd} , while the top x-axis in terms of the volume radius of the primary Roche lobe R_L . The system parameters used to derive the inner and outer radius, R_{in} and R_{out} , are the ones published by Beekman et al. (2000). The parameters R_{wd} and R_L were calculated according to equations (2.82) and (2.4c), respectively, in Warner (2003). A look at Fig. 3 reveals a variety of behaviours. The Balmer

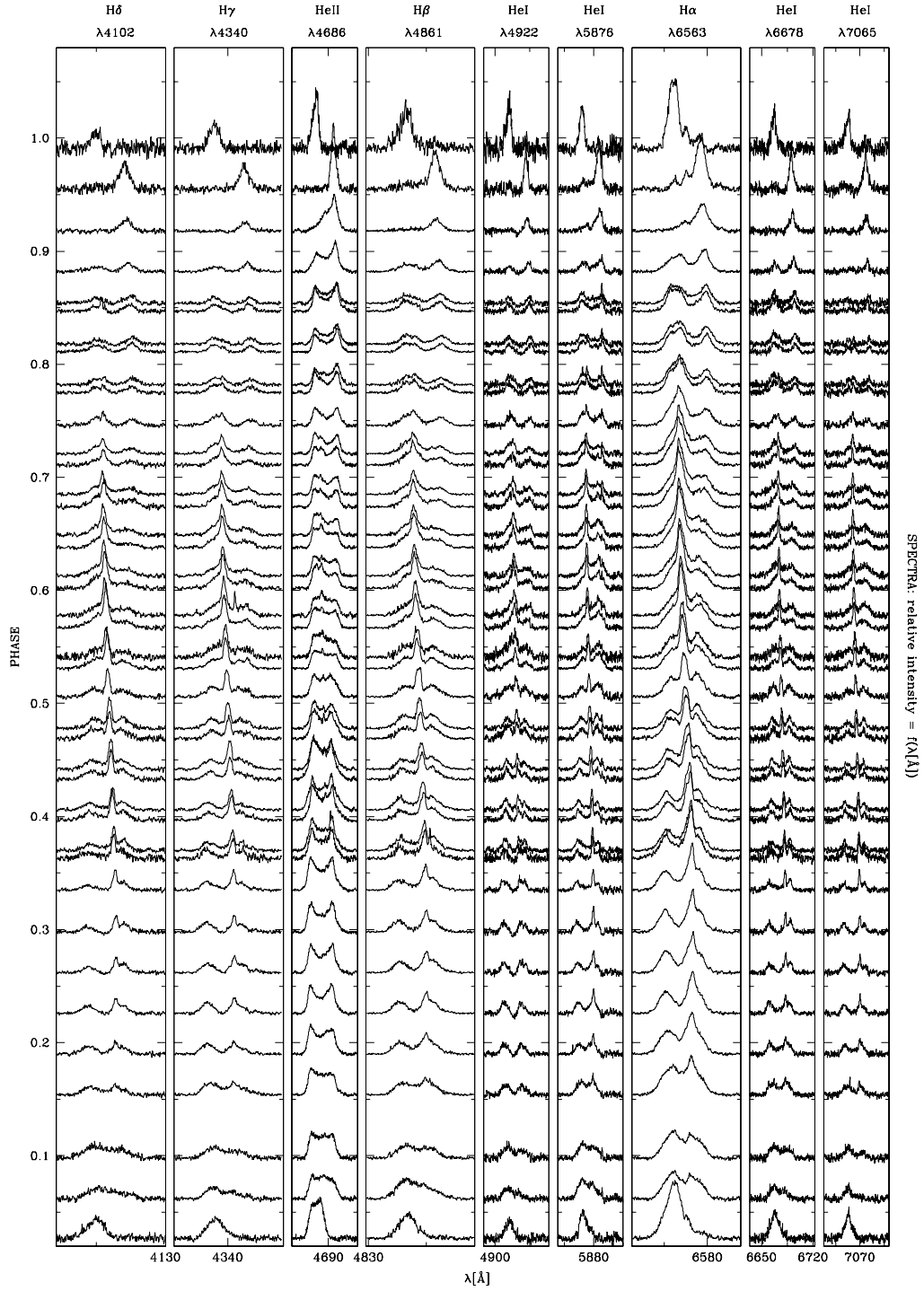


Figure 4: IP Peg trailed spectra of the 9 most prominent lines. The spectra have been continuum normalised and ordered according to phase. The size of the x-axis small and big tickmarks are, in all cases, 35 and 70 $\text{\AA}$, respectively. Note however, that each x-axis has been scaled accordingly, in order for the trailed spectra to be best viewed.

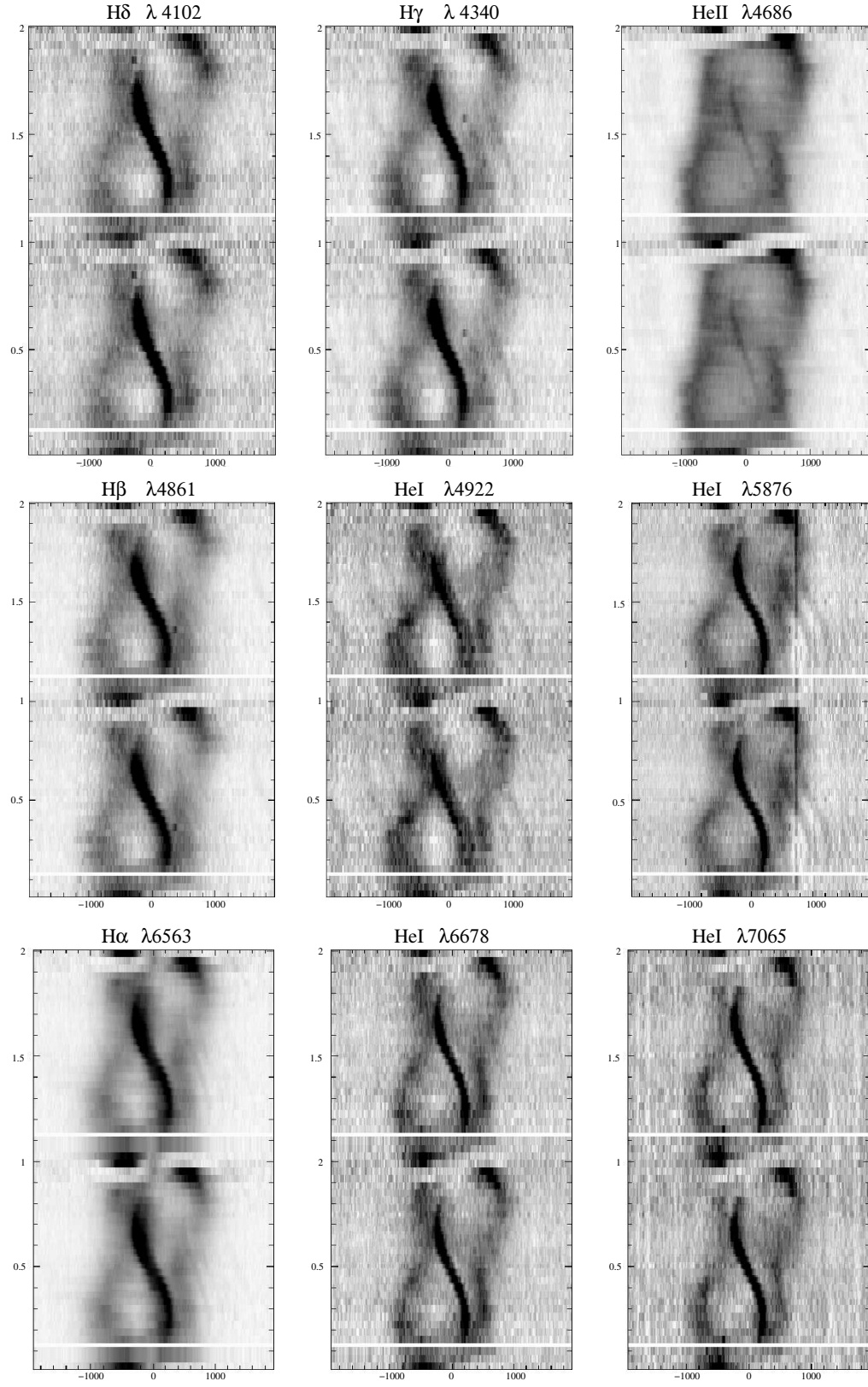


Figure 5: IP Peg spectrograms. The x-axis represents the velocity in km s^{-1} and the y-axis the orbital phase. The grayscale contrast has been adjusted to emphasise between emission (black) and absorption (white).

line emission commences from close to the WD, $R_{\text{in}} \approx 5 \pm 2 R_{\text{wd}}$ and extends up to $R_{\text{out}} \approx 50 R_{\text{wd}}$ for $\text{H}\alpha$. Both R_{in} and R_{out} gradually increase from $\text{H}\beta$ to $\text{H}\alpha$, albeit much more abruptly in R_{out} . The high-excitation line of $\text{HeII } \lambda 4686$ also appears from close to the WD but is constrained to a smaller outer radius $R_{\text{out}} \approx 17 R_{\text{wd}}$ and therefore is the most centrally located line. The $\text{HeI } \lambda 7065, 5876$ lines have a similar $R_{\text{in}} \approx 32 R_{\text{wd}}$ and $R_{\text{out}} \approx 35 R_{\text{wd}}$ but $\text{HeI } \lambda 4922$ appears more centrally located, most probably due to its blending with the FeII line which also causes an asymmetry of the blue-to-red peak not present in the rest of the He lines as seen in Fig. 2. However, $\text{HeI } \lambda 6678$ displays a different behaviour, being more centrally located and having a greater radial extension of $\approx 10 R_{\text{wd}}$, compared to the rest of the HeI lines that are more confined with an extension of $\approx 3 R_{\text{wd}}$. Non-uniform He distribution was also reported by Piche & Szkody (1989), when IP Peg was observed near outburst maximum. The errors of the inner disc radii presented in Fig. 3 are equal to $\approx 2 R_{\text{wd}}$ for the HeI lines and less than $1 R_{\text{wd}}$ for the rest of the emission lines. The errors of the outer disc radii are around $12 R_{\text{wd}}$ for the majority of lines.

After normalising the spectra, and choosing the regions of interest, the trailed spectra of the 9 emission lines were computed. They are ordered according to orbital phase and shown in Fig. 4. In order to enhance the emitting regions, the corresponding spectrograms are presented in Fig. 5, with black corresponding to maximum emission. They have been folded on phases 0–1 and repeated over two cycles. Both figures show characteristics typical of IP Peg: (i) prograde rotation of a disc-like emitting region, with the blue peak being eclipsed prior to the red one, (ii) double-peaked emission lines signature of an AD (iii) highly phase-dependent double-peaked separation, (iv) an S-wave moving in anti-phase with the peaks, signature of chromospheric emission from the secondary star. Additional features, seen around $\approx 1000 \text{ km s}^{-1}$ in the spectrograms of some emission lines, are caused by blendings with weaker emission lines in the red-wings of the principal lines. Each spectrogram has been plotted in a scale that best enhances its emitting regions, and therefore the intensities of the different emission lines should not be directly compared. At this point, the application of the modulation Doppler tomography, presented in the next section, will permit us to gain further insight of the AD structures.

5 Modulation Doppler tomography

Modulation Doppler tomography, developed by Steeghs (2003), is an extension of the classical Doppler tomography which relaxes one of its fundamental axioms, that the flux from any point is constant in time. This allows the mapping of time-dependent emission sources, and this variability information is not considered in standard Doppler tomography studies. Observations show that emission sources can modulate in time and this can happen either because there is anisotropic line-emission or because all sources are not equally visible at all times due to the system's geometry or both. Modulation Doppler tomography gives three output images: one describes the average non-varying flux distribution, and the other two describe the

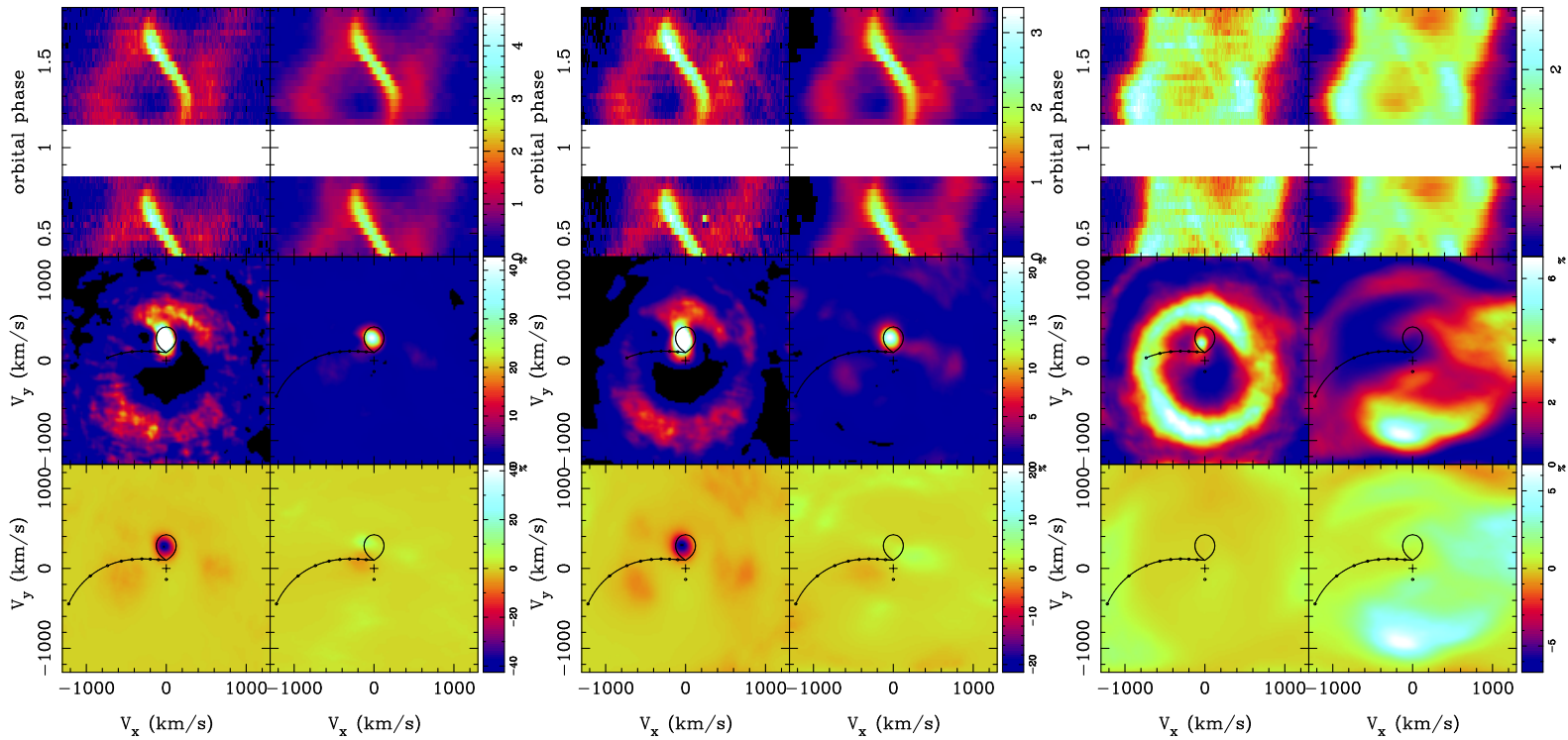


Figure 6: In each panel from top to bottom and left to right: observed data, predicted data, non-variable emission distribution, total variable emission component, cosine and sine terms of the variable component. From left to right: $H\delta$, $H\gamma$, $HeII\ \lambda 4686$.

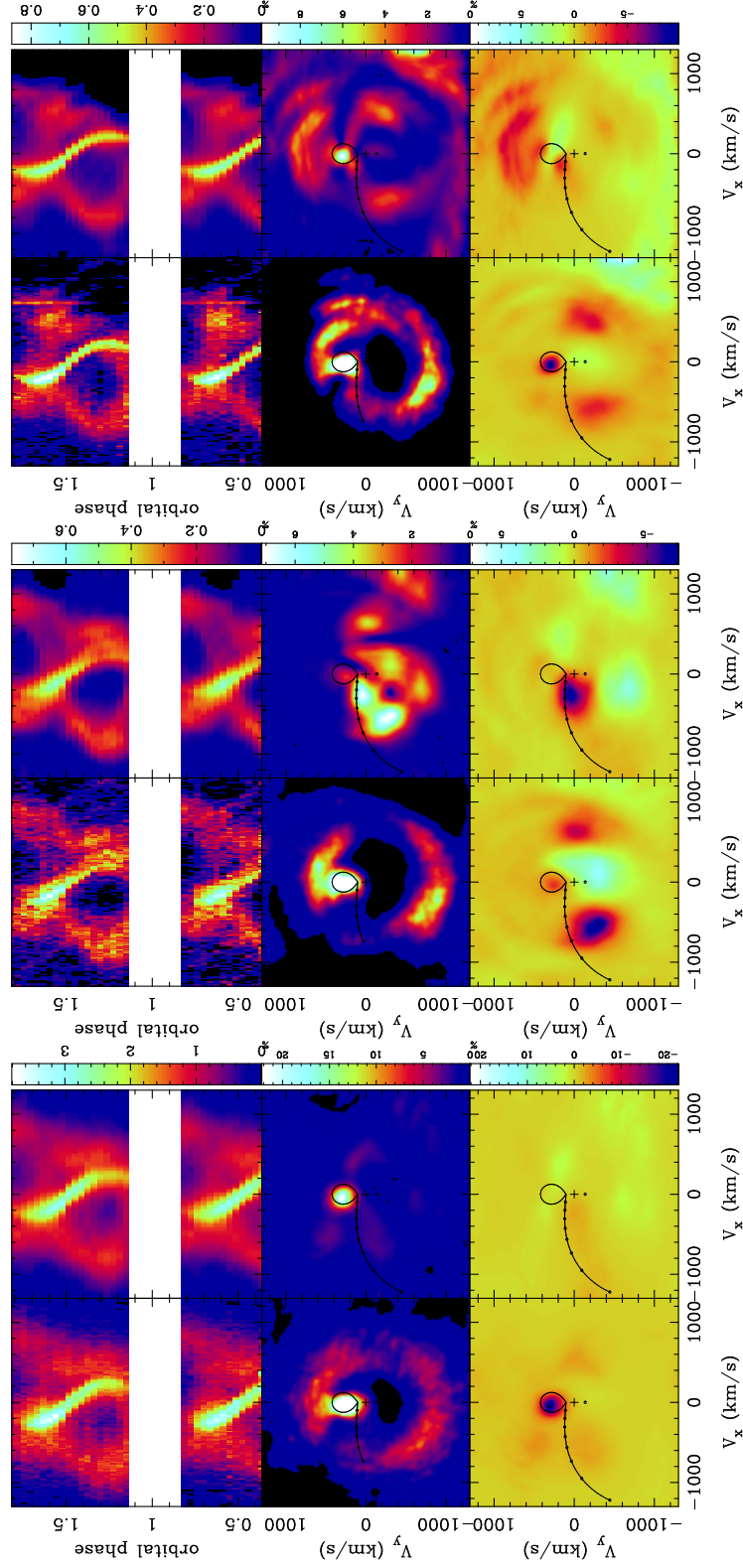


Figure 6: (continued) From left to right: H β , He I $\lambda 4922$, He I $\lambda 5876$.

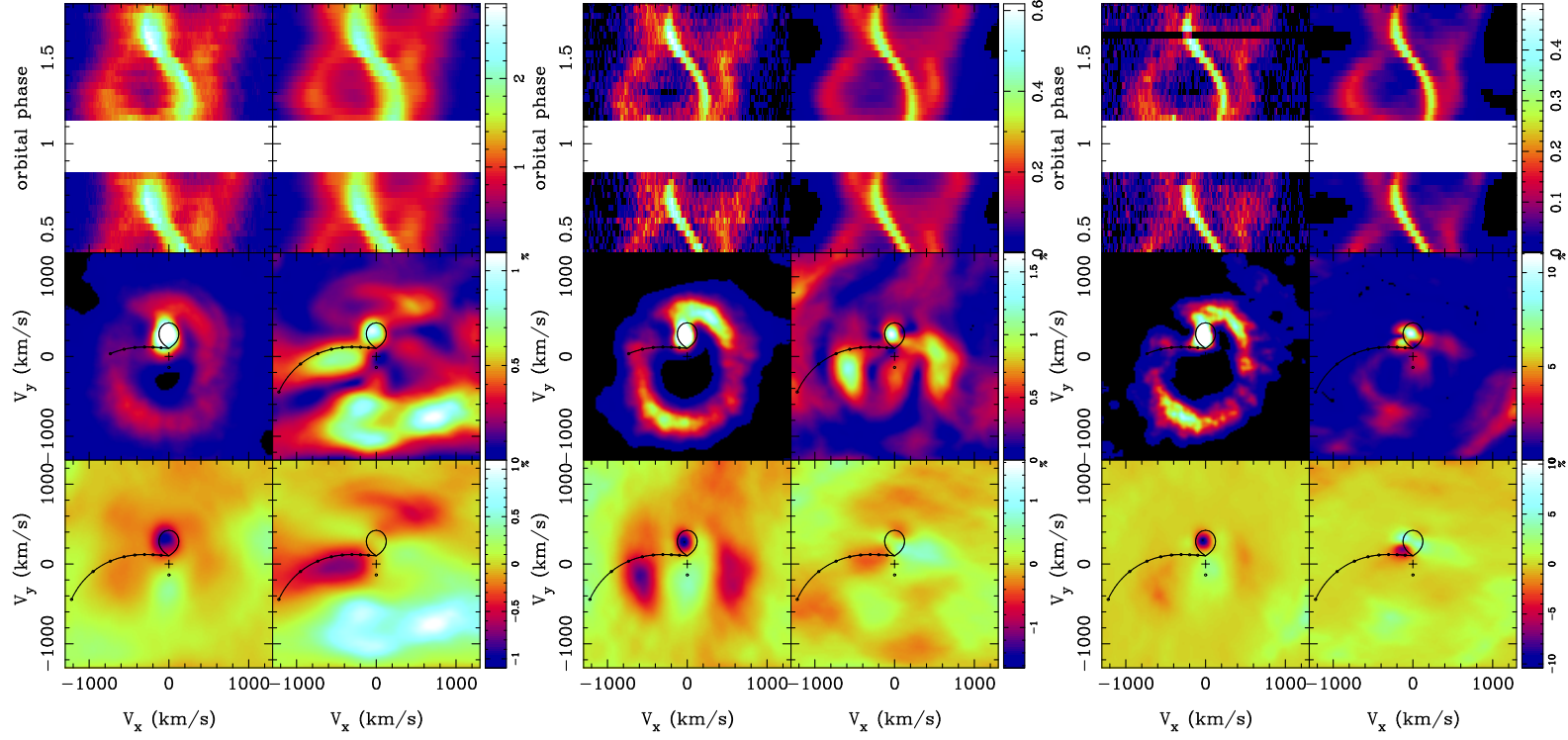


Figure 6: (continued) From left to right: H α , HeI $\lambda 6678$, HeI $\lambda 7065$.

cosine and sine term of the harmonically varying emission sources as a function of orbital phase. For a detailed description of Modulated Doppler tomography, the implementation of the MODMAP code and test reconstructions please refer to Steeghs (2003) and references therein.

The resulting images after running MODMAP, for the 9 most prominent lines, are shown in Fig. 6. From top to bottom and from left to right each panel shows: the observed and predicted data (the latter calculated from the constant and variable components represented by the maps), the non-variable emission distribution and the total variable emission component and finally the individual sine and cosine terms of the variable component. The input and predicted line profiles are plotted in the same scale, reflected by the colour wedge on their right. The average non-varying and total modulated emission maps are plotted in their own scale, with the colour wedge reflecting the scale of the modulated map. The scale is given as a percentage of the average emission to easily depict how strong the modulation is, relative to the constant component. The cosine and sine amplitude maps are also plotted in their own (common) scale, reflected by the colour wedge on their right, again as % of the average emission, in order to ascertain their relative amplitudes. The Roche lobe shape and ballistic stream trajectory have also been computed by the MODMAP code, for the mass ratio and binary separation of IP Peg, and are over-plotted for reference.

The average non-varying-emission maps of all lines show the characteristic spirals plus donor star emission that is well known for IP Peg. The Balmer lines show strong modulation of the donor star, mostly attributed to its cosine term, while $H\alpha$ also shows modulated emission near the stream trajectory. The modulation of the donor star is at a 40% level of the average emission in $H\delta$ and decreases when moving to lower number Balmer lines reaching about 1% for $H\alpha$. There appears to be some modulated disc emission in $H\alpha$ though only at a 1–2% level. No modulated flux, at the signal-to-noise ratio (S/N) of our data, is seen from the spiral arms.

The HeI lines reveal a strong component at the position of the secondary star which is of course modulated. Two disc regions, as viewed in the maps of the cosine term of the variable emission component, appear to show some modulation, though only at the 1–2% level. They also show a component, modulated with the sine term, at the inner Lagrangian point, albeit of low level compared to average emission. The only exception is HeI $\lambda 4922$ (see the cosine map) but one should note that this line is blended with FeII relatively strong in our set of spectra. Again, no significant amount of modulated flux from the spiral arms was needed to reproduce the observed data.

Finally, the high-ionisation line of HeII $\lambda 4686$ also shows the characteristic spiral arms and secondary star emission like the rest of the lines and similarly to the Doppler maps presented by Harlaftis et al. (1999), obtained during another outburst maximum in 1996. The nearby Bowen blend causes some contamination, most likely the cause of the diffuse modulated structure present in the borderline between the 1st ($+V_x, +V_y$) and the 4th ($-V_x, -V_y$) quadrants. It is the only line in which there is a hint of modulated emission at the location of the lower spiral arm and at a 6%

level.

6 Summary & Discussion

This study presents the first time-resolved echelle spectra of IP Peg, one day after the peak of an outburst. Our spectra show the ensemble of the characteristics previously seen in spectroscopic studies of IP Peg, which however, did not cover all lines together. The spectroscopic results, deduced from the average out-of-eclipse spectra are analytically given in Sect. 5 and can be summarised as follows:

- Strong Balmer double-peaked emission arising from the AD. The lines show an equal blue-to-red-peak strength, as opposed to the higher excitation lines that have a stronger blue peak, therefore pointing to a less uniform emission. The Balmer decrement shows a mix of optically thin and thick conditions. The radial extension of the emission originates close to the white dwarf and extends up to $\approx 0.45R_L$ for $H\alpha$. The outer radius of the radial extension gradually increases when moving from $H\delta$ to $H\alpha$.
- Strong double-peaked emission from the high excitation line of HeII $\lambda 4686$ with a stronger blue peak. It is the most centrally located emission with its outer radius reaching $\approx 0.14R_L$.
- Weaker HeI and heavier elements (FeII, TiII and MgII) double-peaked emission. Out of the four HeI lines ($\lambda 7065$, 6678 , 5876 , 4922), HeI $\lambda 6678$ is the most centrally located and shows the greater radial extension.

The trailed spectra and spectrograms, reveal the prograde rotation, the double peaked emission lines arising from the disc, an S-wave component indicative of chromospheric emission from the secondary star and highly phase-dependent double-peaked separation, indicative of a distorted AD resulting in two asymmetric emission components, the spiral waves. All of these features, are present and obvious in the constructed Doppler maps.

We have applied to our data set, the modulation Doppler tomography method, which in addition to the classic Doppler maps reveals if any harmonically varying components are present in the system configuration. This is of great importance in sight of the spiral waves, which nevertheless have not shown any significant harmonically varying emission and therefore we can not claim anisotropic emission. The only line to show a hint of anisotropic emission, is the HeII line, and only at the location of the lower spiral arm and at a 6% level of the average emission, but this could be due to the blending of the line. A reconstruction of the modulation Doppler maps of the HeII $\lambda 4686$ line during another outburst in 1996 (as presented in Harlaftis et al. (1999)) was performed by Steeghs (2003). As in our case, most of the modulated flux was dominated by the sine term. Additionally, the modulated emission was concentrated in two arms (with the lower arm being stronger). The two-armed modulated emission was not attributed to emission originating from the

spiral arms, as there was only partial overlap between the location of the modulated arms and the spiral arms.

The fact that for the majority of lines no modulated emission is seen from the spiral arms does not mean that the spirals do not modulate at all, but that no significant modulated contribution was needed to reproduce the observed data. At higher S/N we may have seen some modulation. But, since we see features of amplitude a few %, this means that our S/N rules out strong anisotropy in disc emission. We could therefore conclude that, if present, the modulated emission of the spiral arms is relatively weak at that our S/N was good enough to put an upper detection limit of any modulated emission at 5–6%.

The Balmer and HeI lines show strong secondary star emission, part of which is modulated. Its level of emission decreases when moving from $H\delta$ to $H\alpha$. As stated in the previous section secondary star emission is naturally modulated with the cosine term. This is normal since irradiation induced emission from the secondary is phase dependent, so it is natural to see a flux variation when the phase is changing. Given equation 3 of Steeghs (2003) and that at phase 0.5–0.7, the face on secondary should emit more, then the cosine term is the dominant component.

All HeI lines show modulated emission in 2 regions of the AD, but are at a very low level. Only in $\lambda 4922$ do they appear more pronounced. These AD modulated regions, seem to appear in the "gaps" between the spiral arms and modulate with the cosine-term. This could be in accordance with the spiral arms shielding the disc areas just behind them and hiding a possible sine-term of disc modulated emission as well. Moreover, the sine-term is expected considering their phasing is in between the spiral arms, and spirals are best viewed at orbital phase 0.5. Geometric shielding caused by the spirals was also proposed from Steeghs (2003) and shielding of the secondary star by the AD was proposed by Watson et al. (2003). Finally, most of the HeI lines show modulated emission near the inner Lagrangian point also covering a small region at the beginning of the stream trajectory. $H\alpha$ though, shows pronounced modulated emission for the bigger part of the stream trajectory.

Last but not least, it has to be mentioned that the blending of some of our lines does not affect the interpretation of the modulation Doppler maps. Blending can be a concern if emission lines overlap in which case multiple lines need to be fitted simultaneously. However for the lines presented in this paper, blending was modest and only affected the line profiles at very high velocities. The advantage of the Doppler technique is that any blending is constrained to map regions at velocities larger than half the separation between the lines. The inner regions are completely unaffected. Since the emission patterns we study in this paper have velocities $< 1000 \text{ km s}^{-1}$, blending was not an issue.

In a future paper it is our aim to study the same dataset, but focusing only to the spiral arms of the emission lines. Tracing the spiral arms of all lines, providing their amplitude and their contribution to the total flux, as well as finding the ratio of the brightness between the two spiral arms can yield valuable results for this unique set of time-resolved spectra.

7 The data

- `fig3.dat`: an ascii file that gives the radial extension of the emission lines as a factor of both the white dwarf radius R_{wd} and the volume radius of the primary Roche lobe R_{L} , as seen in Fig. 3.
- `HJD_phase.dat`: an ascii file giving the heliocentric Julian date and the orbital phase of the 41 spectra.
- `spectra`: a folder that contains,
 - the average out-of-eclipse spectrum of Fig. 2 as a gzipped fits file, named `IP Peg_avr.fits.gz` and
 - the normalised trailed spectra of the 9 most prominent emission lines as shown in Fig. 4. These trailed spectra were the ones used for the reproduction of Fig. 4–Fig. 6. The naming convention is {emission line}_{sequence number}.fits.gz, resulting in 9 (number of emission lines) $\times$ 41 (number of spectra) different files.

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